

The expected detection of dust emission from high-redshift Lyman α galaxies

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ABSTRACT

Recent results have shown that a substantial fraction of high-redshift Lyman α ($\text{Ly}\alpha$) galaxies contain considerable amounts of dust. This implies that $\text{Ly}\alpha$ galaxies are not primordial, as has been thought in the past. However, this dust has not been directly detected in emission; rather it has been inferred based on extinction estimates from rest-frame ultraviolet (UV) and optical observations. This can be tricky, as both dust and old stars redden galactic spectra at the wavelengths used to infer dust. Measuring dust emission directly from these galaxies is thus a more accurate way to estimate the total dust mass, giving us real physical information on the stellar populations and interstellar medium enrichment. New generation instruments, such as the Atacama Large Millimeter Array and Sub-Millimeter Array, should be able to detect dust emission from some of these galaxies in the submillimeter. Using measurements of the UV spectral slopes, we derive far-infrared flux predictions for a sample of 23 $z \geq 4$ $\text{Ly}\alpha$ galaxies. We find that in only a few hours, we can detect dust emission from 39 ± 22 per cent of our $\text{Ly}\alpha$ galaxies. Comparing these results to those found from a sample of 21 Lyman break galaxies (LBGs), we find that LBGs are on average 60 per cent more likely to be detected than $\text{Ly}\alpha$ galaxies, implying that they are more dusty, and thus indicating an evolutionary difference between these objects. These observations will provide better constraints on dust in these galaxies than those derived from their UV and optical fluxes alone. Undeniable proof of dust in these galaxies could explain the larger than expected $\text{Ly}\alpha$ equivalent widths seen in many $\text{Ly}\alpha$ galaxies today.

Key words: galaxies: evolution – galaxies: high-redshift – galaxies: ISM.

1 INTRODUCTION

Lyman α ($\text{Ly}\alpha$) galaxies have been studied in detail for nearly a decade now, and we are only just beginning to learn about the properties of individual objects. Most initial narrow-band studies (e.g. Hu, Cowie & McMahon 1998; Rhoads et al. 2000) observed blank fields with two filters: a narrow-band filter (to capture the $\text{Ly}\alpha$ emission line) and a broad-band filter encompassing this narrow band (used to measure a narrow-band excess). Learning the detailed properties of individual $\text{Ly}\alpha$ galaxies would be difficult with these data. However, these early studies were able to measure

the equivalent widths (EWs) of their objects, calculated from the ratio of narrow-band flux to broad-band flux (Malhotra & Rhoads 2002).

Some of these initial results were surprising, as in many cases the $\text{Ly}\alpha$ EW distribution was higher than one might have expected from stellar population models, as many objects exhibited rest-frame $\text{Ly}\alpha$ EWs over 200 Å (e.g. Kudritzki et al. 2000; Malhotra & Rhoads 2002; Dawson et al. 2004; Shimasaku et al. 2006; Dawson et al. 2007; Finkelstein et al. 2007). A normal stellar population [defined as having a Salpeter (1955) initial mass function (IMF), solar metallicity and a constant star formation rate] has a maximum $\text{Ly}\alpha$ EW of 260 Å at 1 Myr, falling to a constant value of 95 Å at 100 Myr. Thus, while objects with EWs ≥ 200 Å are possible with a normal stellar population, they would all have to be extremely

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young. Looking at a random sample of Ly α galaxies, it would be extraordinarily rare if they all had ages of only a few Myr, thus something is likely happening to skew the EWs higher than they should be.

There are many possible reasons why this could be occurring, and the easiest area to examine is the assumption of a normal stellar population. For example, if a stellar population had subsolar metallicity, it would form stars with hotter stellar photospheres. These stars would then pour out more ionizing photons, which, after interacting with the local interstellar medium (ISM), would in turn create more Ly α photons. Another possibility exists, whereby if a galaxy had a top-heavy IMF, it would create more massive stars than a Salpeter IMF, again resulting in more ionizing photons, and thus more Ly α photons. Both of these scenarios are possible in primitive, dust-free galaxies. Strong Ly α lines also come from the active galactic nuclei (AGN) components of some galaxies, but previous work has shown that most Ly α galaxies detected via narrow-band selection are not AGN (Malhotra et al. 2003; Wang et al. 2004).

Ly α photons are resonantly scattered by any neutral hydrogen present, thus in a homogeneous ISM, they will have to travel much further to escape the galaxy, dramatically increasing the chance of absorption by dust (Meier & Terlevich 1981). However, if the dust is in an inhomogeneous distribution, it is possible that a galaxy could appear to have a strong Ly α EW at later ages.

In this scenario, Ly α can escape unaffected by dust if it is somehow screened from interacting with the dust grains. This can happen if the dust resides in cold, neutral clumps embedded in an ionized interclump medium. This scenario, first proposed by Neufeld (1991), has been examined in detail more recently by Haiman & Spaans (1999) and Hansen & Oh (2006). In this ISM geometry, the Ly α photons get scattered by neutral hydrogen at the surface of these clumps, and thus never see the dust which is embedded deeper inside. However, the continuum photons suffer no such restriction, so they pass un-obstructed into the interior of the clump, with many continuum photons being absorbed (extincted) and scattered (resulting in reddening of the continuum colour). The Ly α photons proceed to bounce around from clump to clump, eventually escaping the galaxy. The net result of this is an increase in the Ly α -continuum flux ratio over that intrinsic to the stellar population. We would thus detect this galaxy with a higher EW, irrespective of its stellar population characteristics.

We believe that this scenario may be responsible for some of the high-EW objects seen in narrow-band studies, and in our previous work (Finkelstein et al. 2009, hereafter F09), we found 10 objects out of a sample of 14 Lyman- α emitting galaxies (LAEs) where this scenario appears to be in effect (one of which was also studied in Finkelstein et al. 2008). Two of these objects are evolved galaxies with large Ly α EWs due to this dust enhancement scenario (the other eight are still young, but have their Ly α EWs enhanced to be larger than their already large value due to dust). The first of these two evolved galaxies was best fit by a 450 Myr, $3.8 \times 10^9 M_{\odot}$ stellar population with a rest-frame EW of $\sim 100 \text{ \AA}$. This relatively high EW at an old age was due to dust, as the best-fitting model had 0.3 mag of dust extinction at 1200 \AA , but this dust was in a geometry such that it did not attenuate Ly α measurably. This resulted in a ~ 35 per cent increase in the Ly α EW over a similar stellar population with no dust. The second of these objects had a most likely age (see F09 for derivation of most likely age) of 500 Myr, a mass of $5 \times 10^{10} M_{\odot}$ and 3.5 mag of dust extinction in a very clumpy geometry, greatly enhancing the Ly α EW.

While these objects are very interesting, and could go a long way towards explaining the apparent discrepancy in EW distributions,

the rest of the objects in the sample had their own surprises. They were all young and lower mass, with ages from 3 to 15 Myr and masses from 1.6×10^8 to $6.1 \times 10^9 M_{\odot}$. These ages and masses are indicative of the traditional belief of Ly α galaxies, where they are very young and possibly primordial in nature. However, all of these 12 LAEs were best fit by stellar populations with considerably non-primordial amounts of dust, with $A_{1200} = 0.5\text{--}4.50$, with many galaxies showing indications of a clumpy ISM.

By fitting their observations to stellar population synthesis models, many other recent studies have derived dust extinction in their Ly α galaxies as well (e.g. Chary, Stern & Eisenhardt 2005; Lai et al. 2007; Pirzkal et al. 2007, hereafter P07; Pentericci et al. 2008). However, these model parameters can be degenerate, as at rest-frame UV and optical wavelengths, the reddening of the continuum colour could be due to both dust and old stars. While observing these objects in the rest-frame optical with the *Spitzer Space Telescope* helps to reduce this degeneracy, data at longer wavelengths could provide much tighter constraints.

When UV and optical photons are absorbed by dust, they are re-emitted in the far-infrared (FIR), resulting in a second observed peak in the spectral energy distribution (SED). Observing these galaxies in the submillimetre (submm) and millimetre (mm) regimes would detect this dust emission directly, which could resolve this red colour degeneracy. The Atacama Large Millimeter Array (ALMA) will be a uniquely sensitive tool to detect these high-redshift galaxies if they have dust, given its excellent observing conditions and large baseline. Thus, obtaining ALMA observations (even non-detections) will provide a much tighter constraint on the amount of dust in these objects than rest-frame UV-optical observations alone. We use ALMA as a measure of future observatory capabilities of detecting dust emission from LAEs. If we can confirm that many Ly α galaxies do in fact have dust, then it will help show that they are not the primitive galaxies that they were once thought to be. On the other hand, if we do not find evidence for dust, it may indicate that a top-heavy IMF or Population III stars are responsible for the strong Ly α EWs.

Where applicable, we assume Benchmark Model cosmology, where $\Omega_m = 0.3$, $\Omega_{\Lambda} = 0.7$ and $H_0 = 0.7$ (cf Spergel et al. 2007). All magnitudes are listed in AB magnitudes (Oke & Gunn 1983). In Section 2, we discuss our sample of Ly α galaxies which we use to make predictions. In Section 3, we present our results, and in Section 4 we make conclusions about the implications of these results.

2 METHOD

2.1 Ly α galaxy sample

We examine samples of LAEs from two recently published studies: F09 and P07. The sample from F09 consists of 14 Ly α galaxies in the Great Observatories Origins Deep Survey (GOODS) *Chandra* Deep Field-South (CDF-S) narrow band selected to be at $z \approx 4.5$. We will use the identifiers from F09, which are based on their detection filters: CHa-1 to 4 (detected in NB656 or H α), CH8-1 and 2 (detected in NB665 or H α + 80) and CS2-1 to 8 (detected in NB673 or [S II]). See F09 for further details on the data reduction and selection process. Seven of the 14 objects have published photometric redshifts putting them near $z \sim 4\text{--}4.5$ (Multiwavelength Southern Infrared Catalog (MUSIC), Grazian et al. 2006; Fireworks, Wuyts et al. 2008). As mentioned in the Introduction, 12 of these objects were best fit by young (3–15 Myr), low-mass ($1.6\text{--}60 \times 10^8 M_{\odot}$) stellar populations with a good amount of dust ($A_{1200} = 0.5\text{--}4.5$),

with the dust in some cases enhancing the Ly α EW. Two objects (CHa-3 and CS2-4) were determined to contain older (450–500 Myr) more evolved populations, with dust extinction attenuating the continuum only due to a clumpy ISM, thus enhancing the Ly α EW.

The sample from P07 consists of nine Ly α galaxies selected in the Hubble Ultra-Deep Field (HUDF; Beckwith et al. 2006), identified by their emission lines with spectra from the Grism ACS Program for Extragalactic Science (GRAPES; Pirzkal et al. 2004; Malhotra et al. 2005; Xu et al. 2007; Rhoads et al. 2008) survey. These objects were selected based on a detectable Ly α emission line in the GRAPES grism data, and broad-band Advanced Camera for Surveys (ACS) photometry consistent with a Lyman break galaxy (P07). In this paper, we will use the object IDs from P07: 631, 712, 4442, 5183, 5225, 6139, 9040, 9340 and 9487. These objects lie in the redshift range $4.00 \leq z \leq 5.76$, with Ly α line fluxes ranging from $17\text{--}60 \times 10^{-18}$ erg s $^{-1}$ cm $^{-2}$ and z' magnitudes from 25.66 to 28.36. Using an exponentially decaying star formation history (SFH), P07 found best-fitting ages of 1–20 Myr and stellar masses from 7×10^6 to $1.8 \times 10^9 M_{\odot}$. Three of the objects were best fit by dust, with 0.05–0.6 mag of extinction in the V band (four objects were best fit with dust using a single-burst SFH and five objects were best fit with dust using a two-burst SFH).

For comparison with our LAEs, we also use the catalogue of spectroscopically confirmed HUDF Lyman Break Galaxies (LBGs) from Hathi, Malhotra & Rhoads (2008; hereafter HMR08; Malhotra et al. 2005; Rhoads et al. 2008). There are 49 LBGs in the catalogue, but for objects at redshifts greater than 4.5, we require J and H band measurements to derive their FIR fluxes, thus we were restricted to using only 21 of the 49 LBGs.

2.2 Dust emission

While we know that photons absorbed by dust are re-emitted in the FIR, the stellar population models used in both F08ab and P07 (Bruzual & Charlot 2003), while computing fluxes as red as 160 μ m rest frame, do not compute dust emission. In order to predict the FIR dust emission from the objects in our sample, we exploit a relation between the UV spectral slope (β ; defined as $f_{\lambda} \propto \lambda^{\beta}$ by Calzetti, Kinney & Storchi-Bergmann 1994) and the IR excess (defined as the ratio of FIR flux to UV flux) detailed in Meurer et al. (1997; hereafter M97). This relation is shown in fig. 1 of M97, with the IR excess increasing exponentially with the spectral slope. This relation is due to dust reddening of the spectral slope (β becoming less negative). Larger quantities of dust will increase the amount of UV extinction, increasing the amount of flux re-emitted in the FIR.

Since we have obtained estimates of dust in our previous studies, we could use our derived extinction values to estimate the FIR flux. However, not all objects are able to be detected in the observed IR, especially at higher redshift. These IR data are necessary to break the aforementioned degeneracy of red colours due to dust or old stars. We aim to predict FIR fluxes independent of stellar population modelling results, with the goal that our method can be applied to objects which only have detections in the optical or NIR.

To use the M97 relation, we first require the UV spectral slope for all objects in our sample, which usually requires spectra. However, in M97, the authors detail a method of obtaining β from the broad-band V and R magnitudes. Unfortunately, this relation is only valid for $z \sim 3$, but HMR08 compute similar relations for $z \sim 4$ and $z \sim 5\text{--}6$. For objects in our sample with $z \leq 4.5$, we computed β via

$$\beta = 5.65 \times (i' - z') - 2, \quad (1)$$

and for objects with $z > 4.5$, we computed β via

$$\beta = 2.56 \times (J - H) - 2. \quad (2)$$

13 of the 23 objects in our sample were calculated to have $\beta \geq -2.5$, indicating that the dust in these galaxies is contributing significantly to their FIR fluxes. Table 1 details the relevant parameters for each object in our sample, including the UV spectral slopes calculated with the above method.

Once β is calculated, we use the relation from M97 to compute the FIR flux from the UV flux (defined in M97 as the flux at 2320 \AA). First, we used the continuum slope calculated in equation (1) or (2), and then from the definition of continuum slope (see above), we calculated the UV flux via

$$F_{\text{UV}} = f_{\lambda,z} \left\{ \frac{[2320(1+z)]^{\beta+1}}{9200^{\beta}} \right\}, \quad (3)$$

where $f_{\lambda,z}$ is the z' flux density in units of f_{λ} and 9200 \AA is the central wavelength of the z' filter. To find $F_{\text{FIR}}/F_{\text{UV}}$, we fit an exponential function to the dashed line in fig. 1 of M97. The equation of this line was

$$\frac{F_{\text{FIR}}}{F_{\text{UV}}} = 66.3e^{1.41\beta} - 1.92. \quad (4)$$

This equation accurately matched M97's fig. 1 over the range of $-2.5 < \beta < 4$. While some objects in our sample did have $\beta < -2.5$, these objects will have a very small $F_{\text{FIR}}/F_{\text{UV}}$ ratio, and thus the emission from their small amount of dust will very likely not be detectable.

In order to obtain the FIR flux densities at a few representative wavelengths, we used equation (10) from M97, which details how the FIR flux is composed of the flux densities at 60 and 100 μ m (in Jy):

$$F_{\text{FIR}} = 1.26 \times 10^{-11} (2.58 f_{60} + f_{100}) \text{ erg cm}^{-2} \text{ s}^{-1}. \quad (5)$$

To independently calculate f_{60} and f_{100} , we require second equation relating these two parameters. We used a template SED from Bendo et al. (2006), who computed the best fit modified blackbody curve to NGC 4631, an edge-on Sd galaxy (a reasonable local analogue to LAEs). They found that a simple blackbody modified by a $\lambda^{-1.2}$ term, with a dust temperature of 28 K, fit their observed data points from ~ 10 to 1000 μ m:

$$B_{\nu}(T) \propto \frac{2h\nu^3}{c^2} \frac{1}{e^{h\nu/kT} - 1} \left(\frac{c}{\nu} \right)^{-1.2}. \quad (6)$$

Adopting an SED of this form and $T = 28$ K, we calculated a ratio of f_{100}/f_{60} of 3.64. Using this in combination with equation (5), we found

$$f_{60} = \frac{F_{\text{FIR}}}{1.26 \times 10^{-11} (2.58 + 3.64)} \text{ Jy}. \quad (7)$$

3 RESULTS

While we know the rest-frame 60- μ m flux for the objects in our sample, we also would like to know how their FIR flux changes with wavelength. To do this, we scaled the Bendo et al. modified blackbody function up to the value of the 60- μ m flux for each object (redshifted to the redshift of the object), so that we could see how their fluxes behaved at all wavelengths in the range of interest. We can then compare the flux densities of these objects at any FIR wavelength to the sensitivities of submm instruments to see if an

Table 1. Properties of LAEs in the Sample^a.

| Name | Redshift | β | f_{UV} (10^{-20} erg cm $^{-2}$ s $^{-1}$) | f_{60} (μ Jy) | f_{100} (μ Jy) | Dust detectable? (Y/N) |
|-------|----------|---------|---|-------------------------|--------------------------|---------------------------|
| CHa-1 | 4.40 | -3.46 | 1.84 | – | – | N |
| CHa-2 | 4.40 | -2.50 | 3.41 | 0.178 | 0.648 | N |
| CHa-3 | 4.40 | -2.90 | 2.76 | – | – | N |
| CHa-4 | 4.40 | -1.21 | 10.46 | 169.1 | 615.4 | Y |
| CH8-1 | 4.47 | -0.35 | 3.97 | 248.1 | 903.2 | Y |
| CH8-2 | 4.47 | -2.67 | 5.13 | – | – | N |
| CS2-1 | 4.53 | -2.22 | 2.96 | 4.733 | 17.23 | N |
| CS2-2 | 4.53 | -0.52 | 2.65 | 129.7 | 472.2 | Y |
| CS2-3 | 4.53 | 1.52 | 8.44 | 7784 | 28333 | Y |
| CS2-4 | 4.53 | -1.14 | 0.65 | 12.05 | 43.87 | Y |
| CS2-5 | 4.53 | -4.70 | 0.32 | – | – | N |
| CS2-6 | 4.53 | -0.10 | 1.37 | 125.3 | 456.1 | Y |
| CS2-7 | 4.53 | -1.65 | 0.44 | 3.296 | 12.00 | N |
| CS2-8 | 4.53 | -2.66 | 0.70 | – | – | N |
| 631 | 4.00 | -2.34 | 2.28 | 1.779 | 6.477 | N |
| 712 | 5.20 | N/A | 0.747 | – | – | N |
| 4442 | 5.76 | -4.94 | 0.145 | – | – | N |
| 5183 | 4.78 | -0.75 | 0.327 | 11.80 | 42.97 | Y |
| 5225 | 5.42 | -3.61 | 3.422 | – | – | N |
| 6139 | 4.88 | -2.03 | 4.191 | 13.63 | 49.61 | Y |
| 9040 | 4.90 | -1.87 | 3.093 | 15.27 | 55.58 | Y |
| 9340 | 4.71 | N/A | 0.521 | – | – | N |
| 9487 | 4.10 | -2.62 | 0.822 | – | – | N |

^aThis table lists some of the relevant properties of the LAEs in our sample. Objects with $\beta < -2.5$ were not able to have their FIR excess derived via the Meurer relation, and thus we do not tabulate their fluxes at 60 and 100 μ m (with such a steep slope, they would likely not have detectable dust emission).

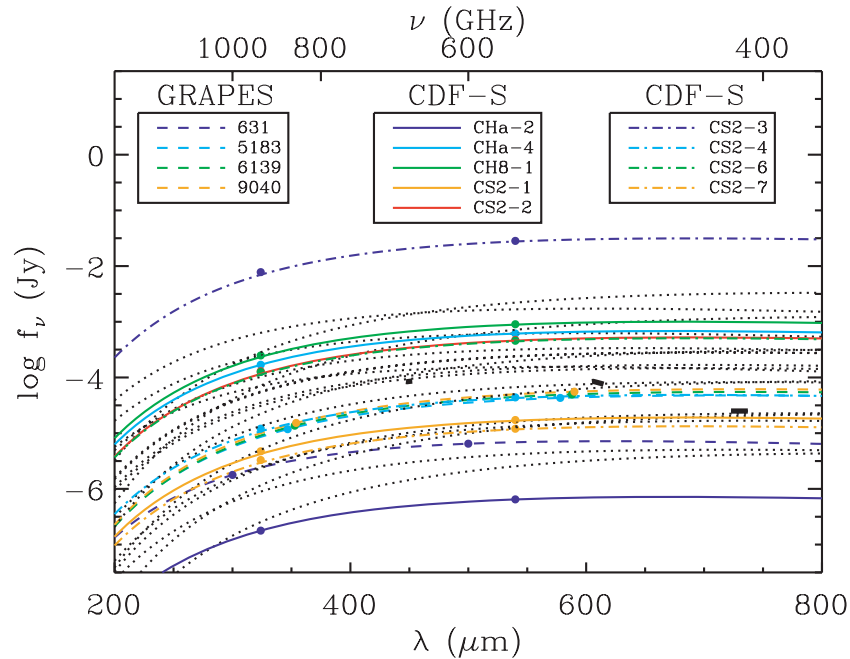


Figure 1. Derived FIR SEDs versus observed wavelength of our objects, with 13 LAEs being shown in the plotted range. The different coloured lines show the LAEs, while the dotted black lines show the SEDs of the sample of LBGs. The curves were derived by scaling the computed 60- μ m flux point to a modified blackbody curve (Bendo et al. 2006). The FIR fluxes of these objects were computed via the relation between the UV spectral slope and FIR excess from Meurer et al. (1997). The thick black lines show the ALMA sensitivities in a 4 h integration, in three windows at 450, 610 and 730 μ m. Out of the whole sample, nine LAEs (13 LBGs) would have their dust emission detected in a 4 h ALMA integration.

object will be detected. With the ALMA sensitivity calculator,¹ we computed ALMA sensitivity curves using the default set of parameters for an integration time of 4 h in bands eight (385–500 GHz) and nine (602–720 GHz). Continuum measurements are possible with ALMA in 8 GHz instantaneous bandpasses, corresponding to a bandpass width of 14 μm at 730 μm (Wilson 2007). We thus chose three windows in which to compare our objects to the ALMA sensitivity curves: 450, 610 and 730 μm . The average sensitivity in an 8 GHz width across these three bandpasses is plotted as thick black lines in Fig. 1. From the relative positions of these sensitivities to the SED curves from our objects in this figure, we can see which objects will have their dust emission detected with ALMA.

3.1 Lyman α galaxies

Out of the 15 objects in the CDF-S sample, Fig. 1 shows that six of them appear to contain enough dust to be detected in the observed submm. These six objects, CHa-4, CH8-1, CS2-2, CS2-3, CS2-4 and CS2-6, contain dust in a range $A_{1200} = 1.0\text{--}4.5$ mag. One of these, CS2-3, which was best fit by $A_{1200} = 4.5$ mag, has the brightest f_{100} data point (28333 μJy) in the whole sample, further proof that our derivation of FIR fluxes is consistent with our model fitting. Objects CHa-1, CHa-3, CH8-2, CS2-5 and CS2-8 have $\beta < -2.5$, thus we did not calculate their FIR fluxes. Objects CHa-2, CS2-1 and CS2-7 do have measurable 60- μm fluxes (0.2–4.7 μJy), but these are likely much too faint to be detectable. It is of note that CS2-4 is only detectable in the 730- μm band.

Three of the GRAPES LAEs which we analysed appear to be bright enough in the FIR to be seen with ALMA, although only in the 730- μm band. These objects are 5183, 6139 and 9040. In P07, object 5183 has a maximum dust extinction² of $A_V = 0.6$ mag (equivalent to $A_{1200} \sim 2.4$), and object 6139 is best fit by 0.4 mag of dust extinction. This is consistent with our results, which imply that the UV light absorbed by this dust is being re-radiated in the FIR. While object 9040 was not best fit by any dust, the authors do acknowledge that its best-fitting model was still a poor fit, although they are confident that it is a LAE at its measured redshift. An important point is that the CDF-S LAEs have a higher detection fraction than the GRAPES LAEs. This is likely due to differences in the sample selection method. The CDF-S sample was selected via ground-based narrow-band imaging (with space-based broad-band photometry). However, the GRAPES sample was selected via space-based grism spectroscopy over a much smaller area (only the HUDF versus the whole CDF-S), thus probing much deeper into the faint end of the luminosity function. As the detectability of an object scales not only with its colour, but also with its continuum brightness, fainter objects will be more difficult to detect.

3.2 Lyman break galaxies

Fig. 1 also plots the derived FIR SEDs of the sample of 21 LBGs. Dust emission from 11 of the LBGs is detectable at 450, 610 and 730 μm , with two additional LBGs being detectable at only 730 μm . In a rough comparison between LBGs and LAEs, 39 per cent of the

23 LAEs are detectable with ALMA, compared with 62 per cent for LBGs. While these numbers may be consistent given the relatively small sample sizes, it could be indicative of a larger difference between the types of objects. Traditionally, LBGs are thought to harbour more evolved and massive stellar populations than those in LAEs. However, from this result, we now see that LBGs may also have a higher dust content than LAEs. This is confirmed by the average β of the two samples, which indicate that LBGs have a redder UV slope with $\langle\beta_{\text{LBG}}\rangle = -1.62$, compared with $\langle\beta_{\text{LAE}}\rangle = -2.03$. A similar effect was seen in Vanzella et al. (2006), who found $\langle\beta_{\text{LBG}}\rangle = -1.7$ and $\langle\beta_{\text{LAE}}\rangle = -2.1$, as well as in Pentericci et al. (2007), who found that LBGs were dustier than LAEs. This is interesting, given that in F09 we found that our sample of LAEs had much higher dust extinction rates than we would have thought. Thus, while LAEs appear to have some amount of processed ISM, the ISM of LBGs appears to be much more enriched, resulting in higher dust emission in the FIR.

3.3 Robustness of results

We have chosen a value of $f_{100}/f_{60} = 3.64$ from Bendo et al. (2006), which corresponds to a dust temperature of 28 K. However, there is a large uncertainty in the FIR colour. Larger values of f_{100}/f_{60} imply a cooler ISM (thus, a cooler dust temperature), while smaller values imply a warm ISM, more consistent with compact star formation (Dale et al. 2001). Dale et al. (2001) and Dale & Helou (2002) compute model FIR SEDs assuming a large range in FIR colour, from $0.8 < f_{100}/f_{60} < 3.3$. Since LAEs are expected to be undergoing significant bouts of star formation, and they are quite compact (e.g. P07), it is interesting to see how our derived FIR SEDs would change if we assume a warmer value of f_{100}/f_{60} . The observations presented in Dale & Helou (2002) show that the number of objects drops off considerably below $f_{100}/f_{60} = 1$, so we assumed $f_{100} = f_{60}$, and redid our calculations. Fig. 2 shows these results for CS2-2, which was one of the fainter LAEs detected in all ALMA bands. Since warmer FIR SEDs will result in fainter fluxes in the ALMA observable range, this plot can give us an idea of the limits on our predicted fluxes.

We find that although the shape of the SED has changed as expected, this particular object would still be detected. However, since the peak of the dust emission has moved towards shorter wavelengths, objects which were only detected in the 730- μm band may be less likely to be detected if they have a warmer SED. As another illustration of our uncertainties, we have plotted the 1σ errors on our original predicted 60- and 100- μm fluxes, obtained from the photometric i' and z' errors. We also plot error bars derived from the uncertainty in the M97 relation, as described in Meurer, Heckman & Calzetti [1999; 0.04 dex uncertainty in $\log(F_{\text{FIR}}/F_{\text{UV}})]$. These errors are much smaller than the photometric errors. However, it is worth deeper investigation into the extent which the photometric errors can alter our results.

Using equations (1)–(6), we can calculate the dust emission from any object given a z' magnitude and a colour ($i' - z'$ for redshifts less than 4.5 and $J - H$ for redshifts greater than 4.5). Figs 3 and 4 show these results at 730 μm , with contours showing the flux levels for any colour–magnitude combination, and one contour specifically showing the ALMA 4 h sensitivity at 730 μm . Fig. 3 plots all objects in our whole LAE + LBG sample with $z \leq 4.53$. The mean redshift of this subsample is 4.33, so the flux curves were derived for this redshift. Fig. 4 plots all objects with $z > 4.53$, with the flux curves derived for the mean redshift of this subsample, which was 5.26.

¹ Available online at <http://www.eso.org/projects/alma/science/bin/sensitivity.html>.

² This value of dust extinction is the maximum amount of extinction present in this object, because the models were fit to this object using upper limit constraints on the rest-frame optical fluxes.

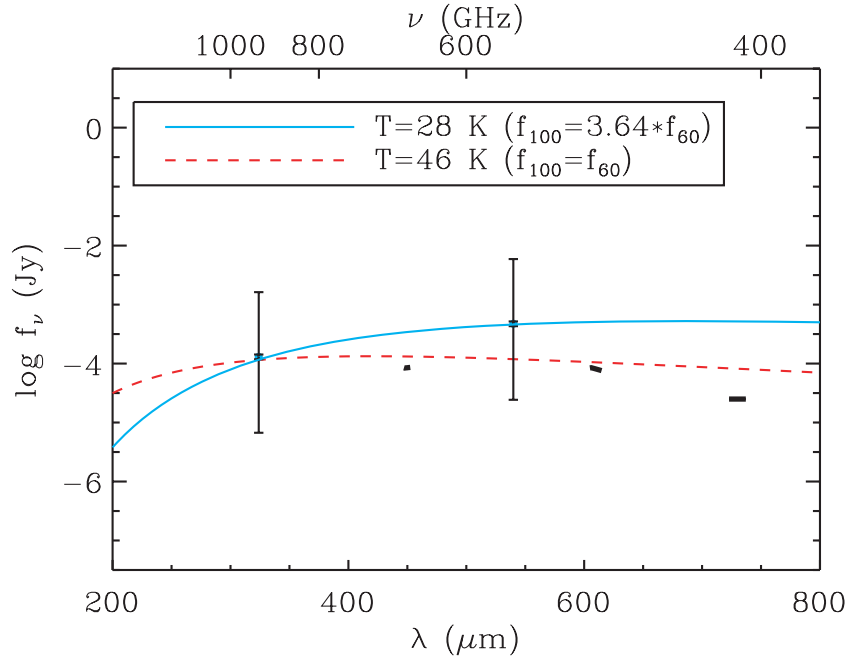


Figure 2. Same plane as Fig. 1. The blue line plots the predicted rest-frame FIR SED for CS2-2. The literature shows that the temperature of the SED depends on whether the star formation proceeds in compact or widespread environments. We thus show (in red) our predictions if we assumed a warmer SED, with $T = 46$ K. The change in temperature does not affect the detection of this object, although it does move the peak of the dust emission towards lower wavelengths. This can affect our results when an object is only detected in the reddest ALMA band. We also show the derived errors on our original SED, obtained from the i' and z' photometric errors. The tiny error bars inside the photometric errors are the errors derived from the scatter in the M97 relation (Meurer et al. 1999).

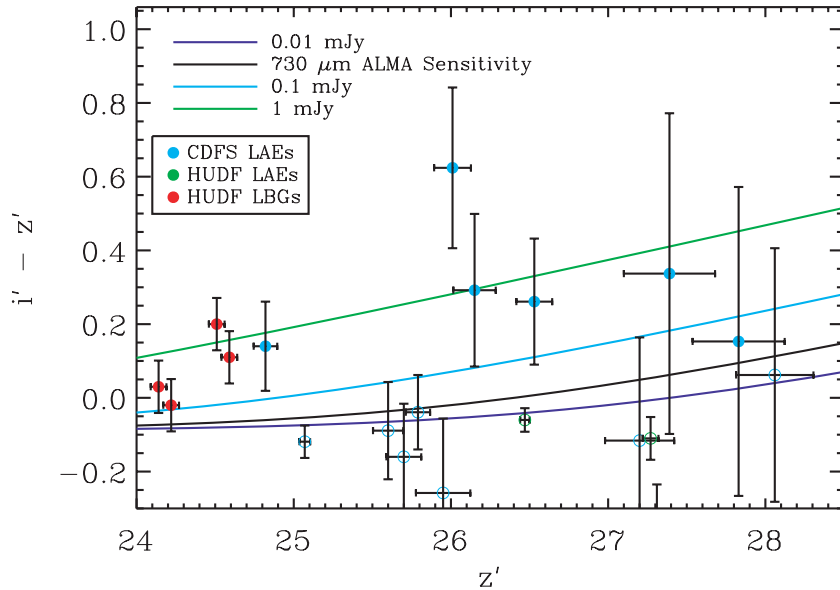


Figure 3. A colour–magnitude diagram, depicting the position of all objects with $z \leq 4.53$. Filled circles represent objects which would be detected with ALMA at $730 \mu\text{m}$ (and possibly also at 450 and $610 \mu\text{m}$), with open circles representing those objects which would be undetected (based on Fig. 1). Blue circles represent the CDFS LAEs, while green denotes the HUDF LAEs and red the HUDF LBGs. The contours show the observed fluxes at $730 \mu\text{m}$ for a computed grid of magnitudes and colours ($i' - z'$). The black contour shows the 4 h ALMA $730 \mu\text{m}$ sensitivity limit. The contours were all computed for $z = 4.33$, which is the mean redshift of all objects in our sample at $z \leq 4.53$. The error bars on the data points are the 1σ errors.

The different coloured circles represent the three different samples (CDF-S LAEs, HUDF LAEs and HUDF LBGs), with filled circles representing objects we found that would be detectable with ALMA via Fig. 1. All filled circles lie above the ALMA 4 h sensitivity contour (black line) showing that Figs 1, 3 and 4 are mutually

consistent, although the three detectable HUDF LAEs do lie less than 1σ above the ALMA 4 h sensitivity contour. Using these figures, one could then take the colour, magnitude and rough redshift of any object, and determine if the dust emission could be detected by seeing if the points lie above the ALMA sensitivity contour.

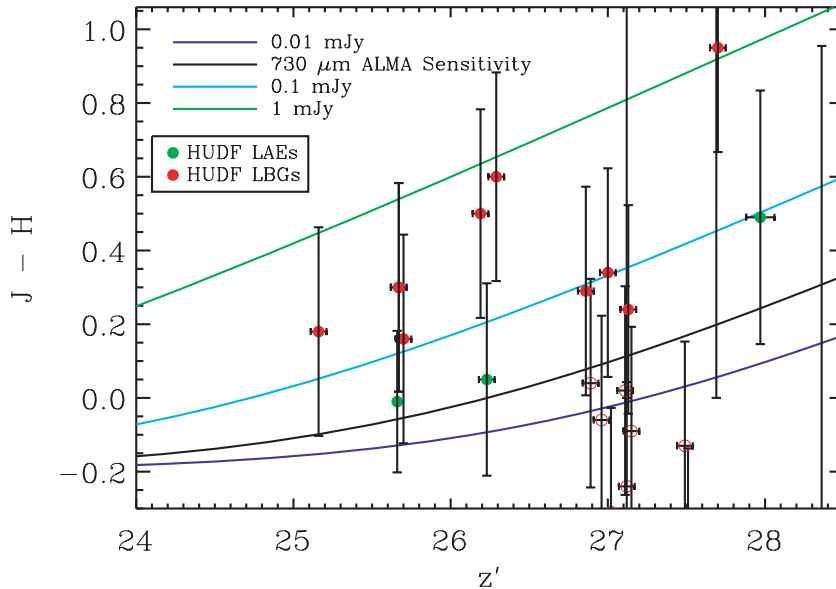


Figure 4. The same as Fig. 3, for all objects with $z > 4.53$, thus the colour is $J - H$. The contours were all computed for $z = 5.26$, which is the mean redshift of all objects in our sample at $z > 4.53$.

Figs 3 and 4 allow us to examine the robustness of our results, as we have visualized the 1σ errors on our objects fluxes. We can thus see how close our objects are to the ALMA sensitivity limit as a function of their photometric errors. While 39 per cent of the LAEs will have their dust emission detectable, only 17 per cent (4/23) of them lie more than 1σ from the ALMA sensitivity contour. Likewise, for LBGs, only 33 per cent lie more than 1σ from the ALMA 4 h sensitivity contour. However, some objects which lie below the ALMA contour are less than 1σ away. Thus, we can define a maximum percentage of objects with detectable dust emission, including not only those above the line, but also those that are below but within 1σ . We find this to be 61 per cent for LAEs and 81 per cent for LBGs. This allows us to put an error bar on our LAE detection fraction, of 39 ± 22 per cent (likewise 62^{+19}_{-29} per cent for LBGs).

While the FIR- β relation was derived at low redshift, both Seibert, Heckman & Meurer (2002) and Reddy & Steidel (2004) find that it also applies at high redshift. In addition, Teplitz et al. (2006) found that when they compared FUV to *Spitzer* observations, the results were in general agreement with the FIR- β relation. Chapman et al. (2005) have investigated this relation using a sample of submm selected galaxies at $z \sim 2$. The authors conclude that the bolometric luminosity of submm galaxies is underestimated by two orders of magnitude when derived from their UV properties. This indicates that our derived FIR fluxes may be lower limits, meaning that if anything we could be underestimating the detection fraction.

However, Reddy et al. (2006) found that the M97 relation may not apply to all star-forming galaxies at $z \sim 2$. They find that galaxies with derived ages < 100 Myr exhibit a smaller FIR excess than that predicted by the M97 relation, indicating that these young galaxies have redder UV colours for a given amount of dust extinction than older galaxies. This is of concern as most of our LAEs have very young derived stellar population ages. In addition, Baker et al. (2001) find that the M97 relation also overpredicts the FIR flux for the gravitationally lensed LBG cB58. However, there are differences between these objects and LAEs, as they are at a lower redshift, and they are continuum selected. We thus cannot know if the same age

dependence will apply to LAEs until we obtain a statistical sample of rest-frame FIR detections of LAEs. In addition, the Reddy et al. figure which illustrates the age discrepancy does not have error estimates on their young object points, thus it is hard to discern how significant this effect is. Baker et al. do show error bars, and these errors imply that their observations are only off the M97 relation by $\sim 1\sigma$.

If young objects do have systematically lower $F_{\text{FIR}}/F_{\text{UV}}$ ratios, it would be worrying, as it could imply that we are overpredicting the FIR fluxes of our LAEs by as much as an order of magnitude. It is thus prudent to apply our results to any relevant observations to see if our FIR flux predictions are accurate. Although very few galaxies with Ly α emission have been detected in their rest-frame FIR, Capak et al. (2008) have recently published a study of a spectroscopically confirmed mm starburst galaxy, which does exhibit Ly α emission. They publish the observed SED of this galaxy from X-rays to radio wavelengths. Using the published optical magnitudes, we can compute β , and predict the FIR flux. Doing this, we calculate a predicted 554 (rest 100) μm flux of 2637^{+5049}_{-1875} μJy . However, the closest point in the observed SED is at 1.1 mm, at 4800 ± 1500 μJy . The flux of our derived SED at 1.1 mm is 2054^{+3932}_{-1460} μJy , thus they are consistent within 1σ . These observations give evidence that our predictions do match observations. We show the SED from Capak et al. along with our predictions for this object in Fig. 5. Additionally, we also show the V , i' , z' and 3.6 μm fluxes for CH8-1, which was one of our redder LAEs. These points have been scaled to match the i' flux of the Capak et al. object. We can see that the $i' - z'$ colour is somewhat redder than the Capak et al. object (0.3 versus 0.15), which indicates that CH8-1 is dustier, explaining its higher expected FIR flux. However, as it is intrinsically fainter (the scaling factor was ~ 17), it would be much harder to detect with current instrumentation.

Although we have assumed a reasonable SED shape, it is possible that the FIR SED could be much different. Houck et al. (2004) discuss the low-metallicity blue compact dwarf galaxy SBS 0335-052, which has a FIR spectrum much different than typical starburst galaxies, peaking at ~ 28 μm , indicating much warmer dust,

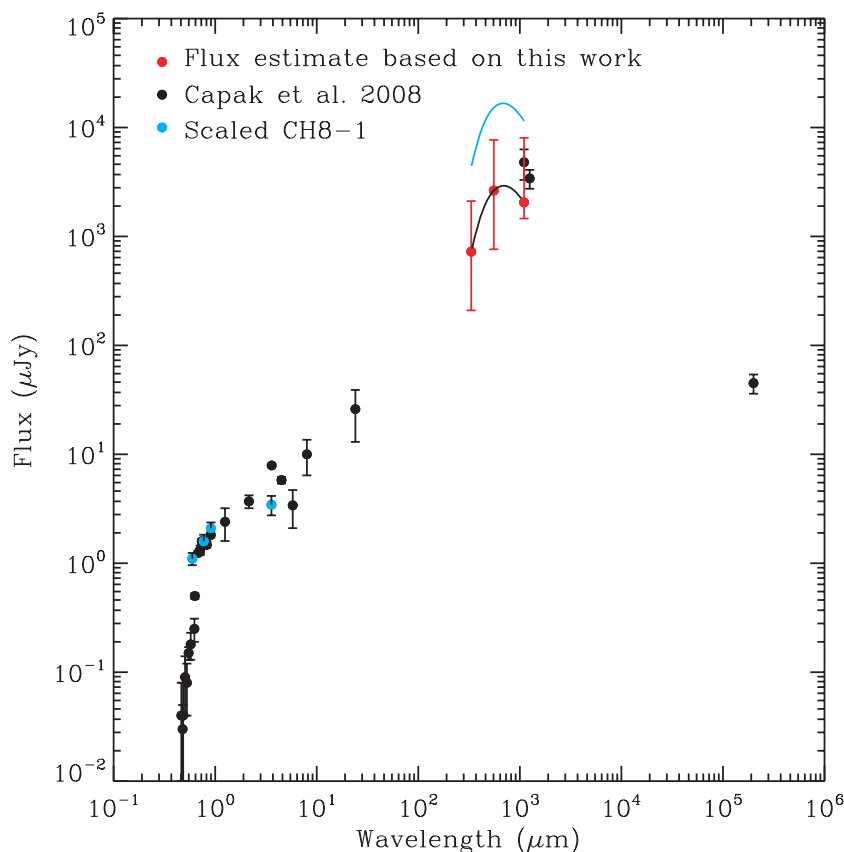


Figure 5. In this figure, we plot the SED from Capak et al. (2008) for a mm detected Ly α emitting starburst galaxy, spectroscopically confirmed to lie at $z = 4.547$. The predicted FIR fluxes match up well with the observed SED, providing further evidence that the M97 relation is applicable at high redshifts. The blue points denote the V , i' , z' and $3.6 \mu\text{m}$ fluxes from CH8-1. These points have been scaled to match the i' flux of the Capak et al. object. We can see that the $i' - z'$ colour is redder than the Capak et al. object, which indicates that CH8-1 is dustier, which we can see as its (scaled) predicted FIR flux is brighter than that of the Capak et al. object. However, this object is fainter by about a factor of ~ 17 , thus it would be much harder to detect with current instrumentation.

possibly due to the lower metallicity. High-redshift LAEs may be expected to have a very low metallicity, thus it is worth mentioning that if any of our objects had this type of SED, we could be vastly overestimating their FIR flux. We have thus investigated the dependence of our results on these derived FIR fluxes. To do this, we reduced our FIR flux to be 50 per cent of its derived value, and computed the ALMA detection fraction of both types of objects. We calculated a detection fraction of 26 per cent for LAEs and 62 per cent for LBGs (57 and 76 per cent for LAEs and LBGs, respectively, within 1σ of the ALMA contour). Thus, while we believe that the M97 relation is applicable to our objects, we find that our conclusions are robust to a factor of 2 variation in the FIR flux.

4 CONCLUSIONS

Directly observing dust emission from Ly α galaxies will provide undeniable proof that these earliest of galaxies have undergone some enrichment of their ISM. While these observations are currently difficult, we have been able to predict whether a sample of 23 LAEs would be detected with a future submm observatory such as ALMA based on their optical and NIR fluxes. Exploiting a relation from M97, which relates the FIR excess to the UV continuum slope, we have derived the FIR fluxes from these LAEs, which will be primarily due to dust emission. By comparing these fluxes to the sensitivity of ALMA at 450, 610 and 730 μm , we have determined

that 39 per cent (9/23) of our sample of LAEs should be detectable based on their dust emission.

We perform the same analysis on a sample of 21 Lyman Break Galaxies, finding that 62 per cent (13/21) of the LBGs are detectable with ALMA, ~ 60 per cent higher than the detection fraction of LAEs. This may be a function of our small sample size, or it may be the result of a true dichotomy in the dust contents of LAEs and LBGs, providing further proof that LBGs are a more evolved class of high- z galaxy. In order to generalize our results, we computed the FIR fluxes for a wide range of continuum magnitudes and colours, comparing these fluxes to the ALMA 4 h sensitivity limit. Thus, knowing the continuum magnitude and colour of any object, one can then predict whether its dust emission will be detectable.

While stellar population modelling gives some insight on the amount of dust in high-redshift galaxies, future submm instruments will allow us to directly observe the dust emission from these objects, providing a more complete census of dust. From our analysis of LAEs and LBGs in the HUDF, we calculate a number density of ALMA detectable LAEs and LBGs of 0.26 and 1.13 arcmin^{-2} , respectively. Thus, even in a small survey area, one could expect to detect a number of high-redshift objects with ALMA in a reasonable amount of time. In our previous analysis of LAEs, the addition of the clumpiness parameter allows strong Ly α flux to an arbitrarily old and dusty galaxy, thus these ALMA observations will help to constrain the amount of dust, and thus the level of clumpiness.

These results will have considerable impact on several areas of galaxy evolution, specifically on the star formation rate (e.g. Giavalisco, Koratkar & Calzetti 1996), which, being derived from UV observations, is critically dependent on accurate measurements of the dust content. Combining ALMA with UV observations, we will have a complete map of star formation in high-redshift galaxies (i.e. Gao 2008). Studying the submm emission of these objects will also allow calculation of the total dust mass, which is something one cannot calculate from UV-optical data alone. Along with learning about the total amount of dust in these objects, ALMA will give us the opportunity to investigate the distribution of dust due to its excellent resolution. At 730 GHz, ALMA can achieve resolutions down to $0.01''$ pixel⁻¹. Combining ALMA observations with high-resolution *Hubble Space Telescope* (*HST*) data will allow us to investigate dust on local scales at high redshift for the first time. Strong submm detections combined with strong Ly α emission could indicate that dust enhancement of the Ly α EW is widespread.

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